

Self-similar galaxy dynamics below the de Sitter scale of acceleration

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Tensions shimmering with Physics beyond Λ CDM?

H_0 tension problem (3.8σ , Riess et al. 2017)

de Sitter is unstable?

Λ CDM ruled out in the future,

Λ CDM cannot hold to all orders today



Ó Colgáin, van Putten, Yavartanoo, arXiv:1807.07451 (2018)

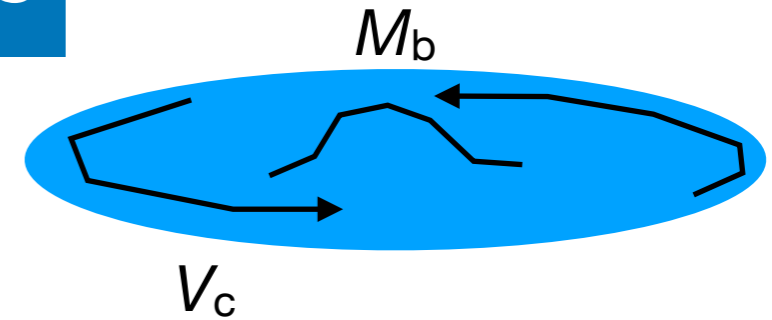
Λ CDM galaxies vs SPARC data

6σ gap about $a_{\text{ds}} = cH$

Galaxy dynamics tracing H ?

this talk

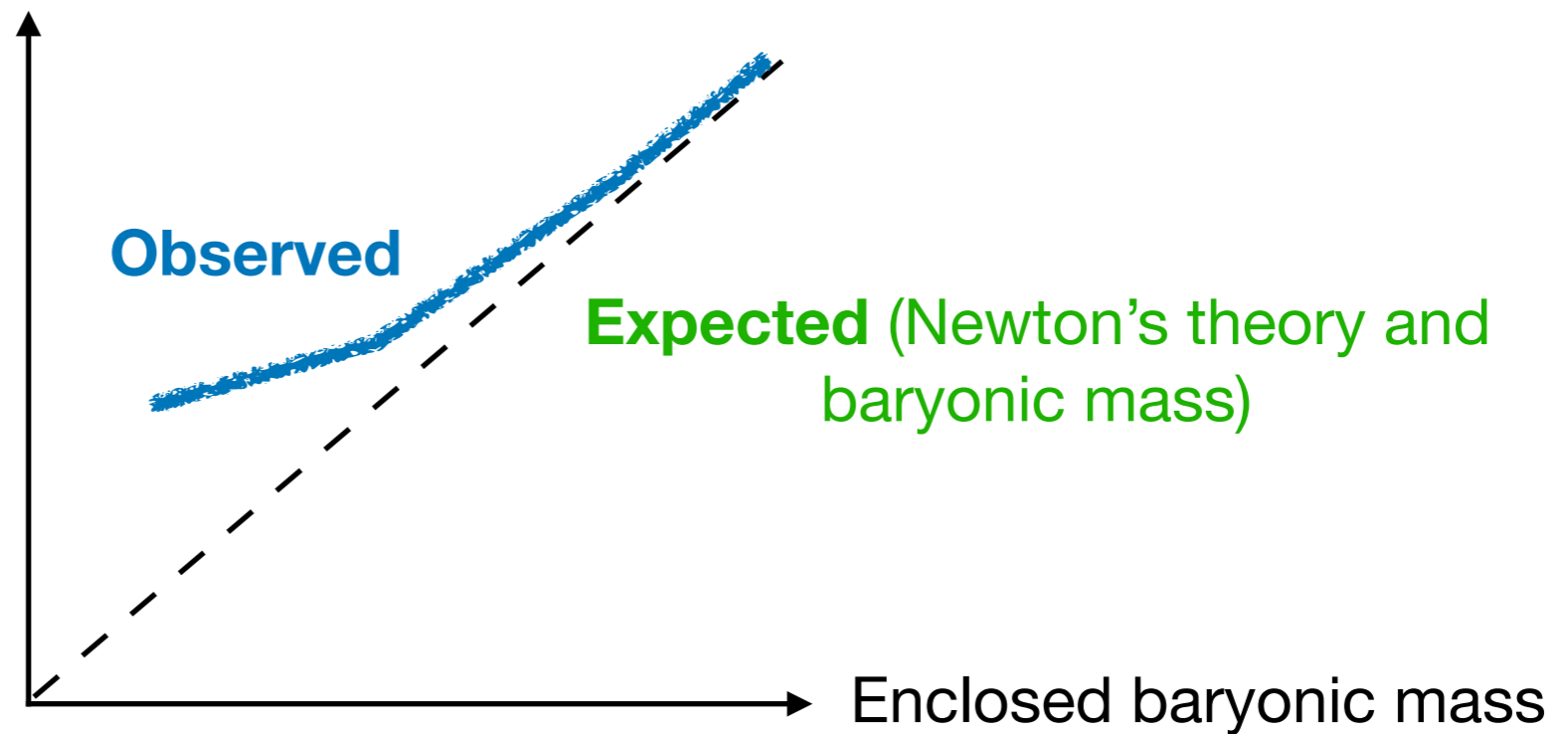
"Missing mass" in spiral galaxies



Spectroscopy

Observed/Expected
radial acceleration

$$\frac{\alpha}{a_N} = \frac{V_c^2 / r}{V_b^2 / r}$$



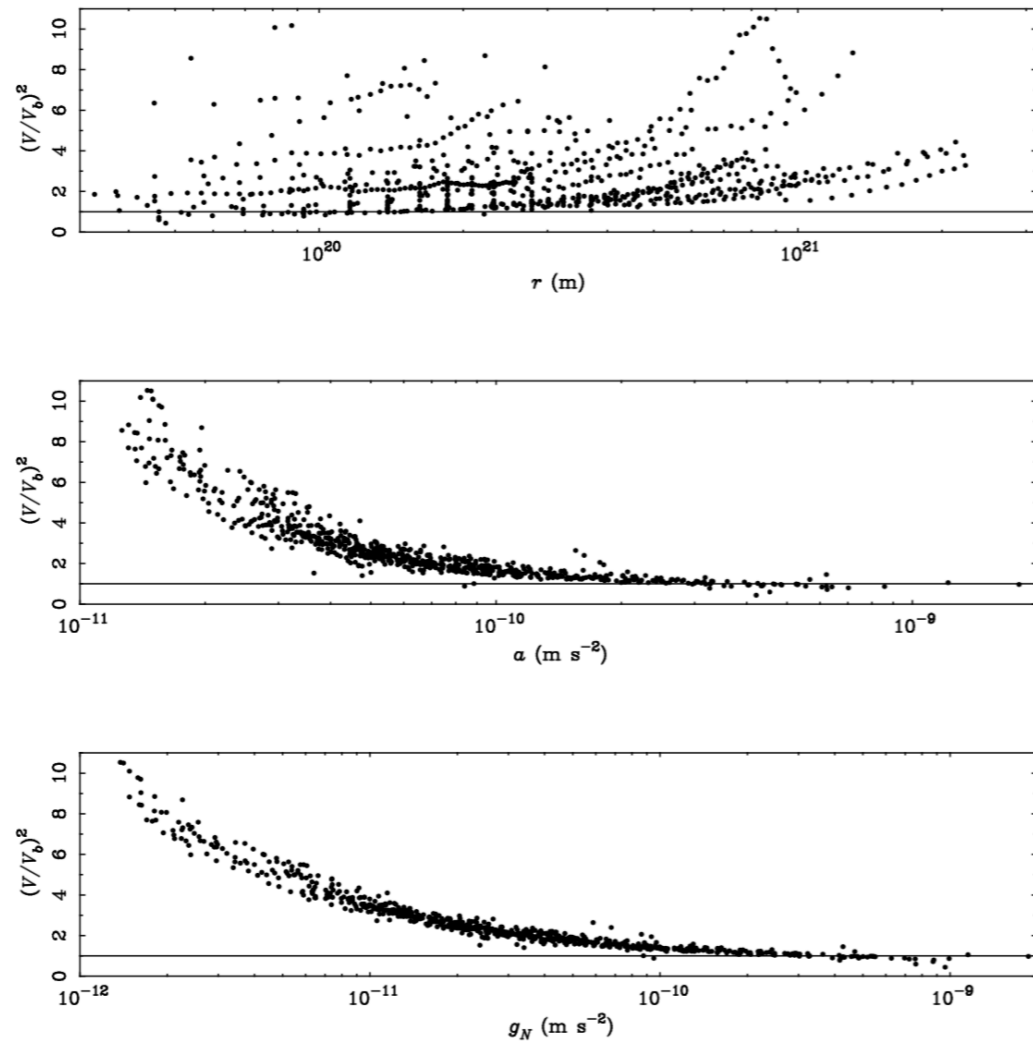
Photometry

$$a_N \equiv -\frac{GM_b}{r^2}$$

Famae, B., & McGaugh, S.S., 2012, LRR, 15, 10

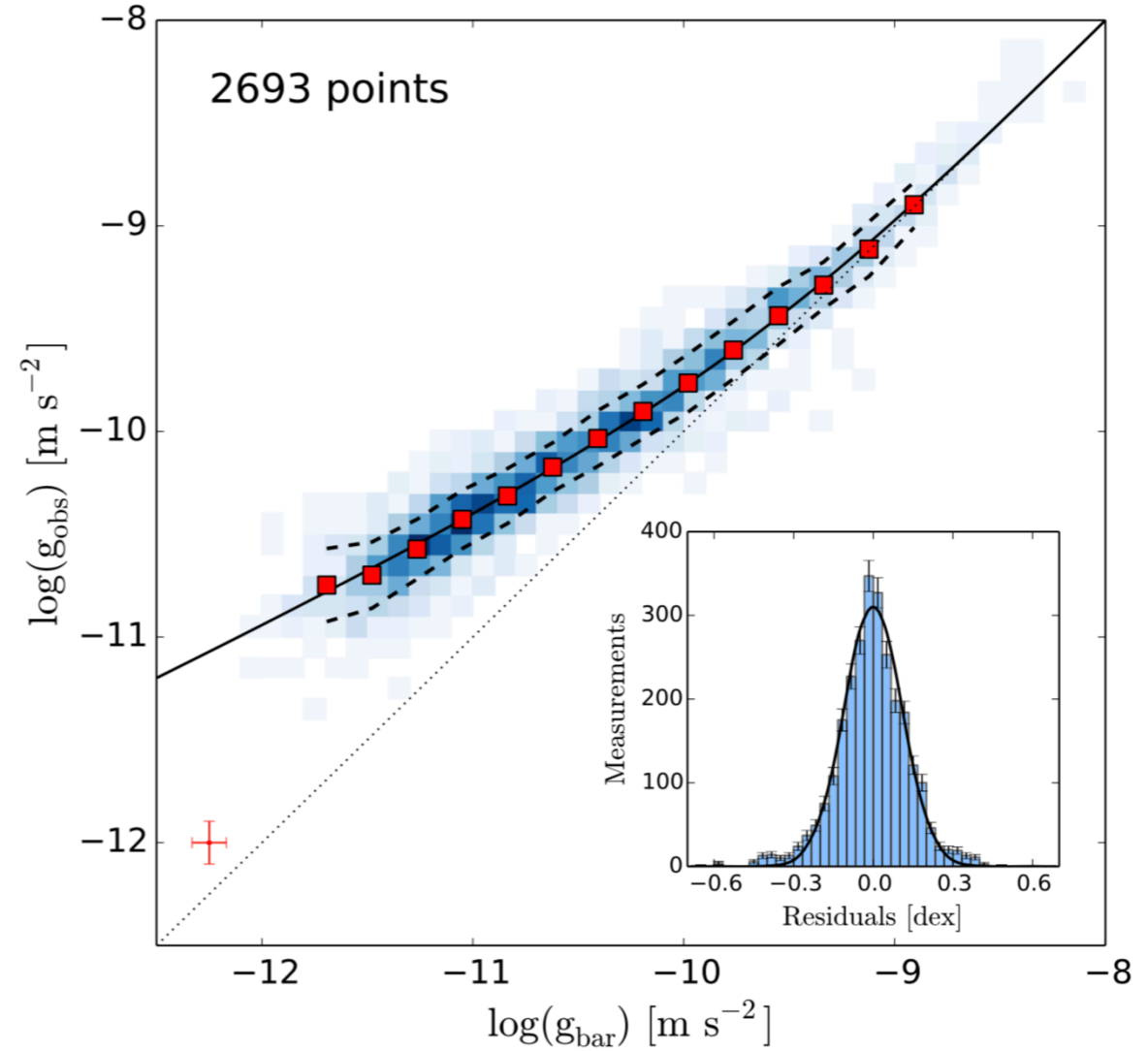
McGaugh, S.S., Lelli, F., & Shombert, J., 2016, PRL, 117, 201101

SPARC data



Famae, B., & McGaugh, S.S., 2012, LRR, 15, 10

SPARC: Spitzer Photometry and Accurate Rotation curves Catalogue ($z=0$)



McGaugh, S.S., Lelli, F., & Shombert, J., 2016, PRL, 117, 201101

Normalize to a_{dS}

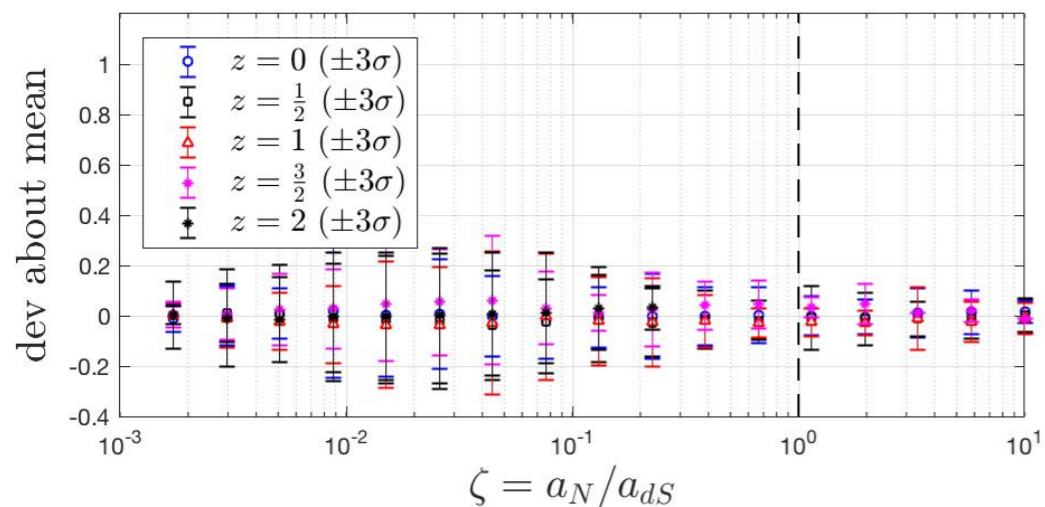
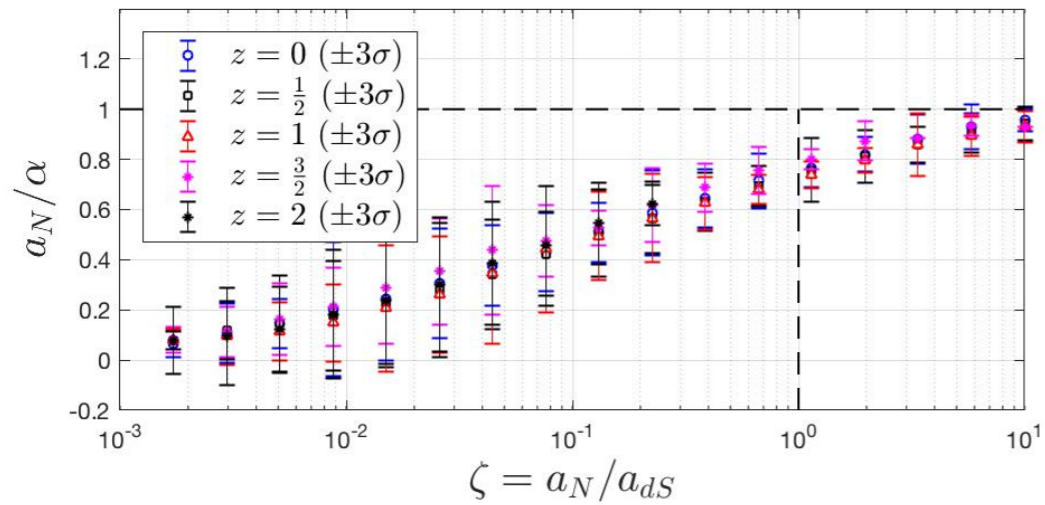
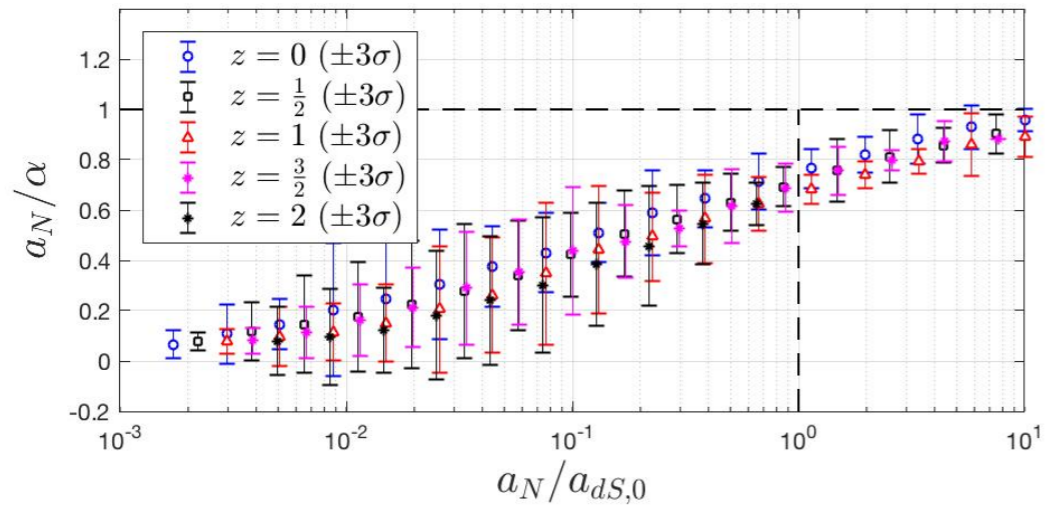
Photometry

$$\zeta = \frac{a_N}{a_{dS}}, \quad a_{dS} \equiv cH$$

Spectroscopy

Radial acceleration⁻¹ $\frac{a_N}{\alpha}$

Λ CDM galaxy models in MUGS2



McMaster Unbiased Galaxy Simulations 2

Keller, B.W., Wadsley, J., Benincasa, S.M., & Couchman, H.M.P., 2014, MNRAS, 442, 3013

Keller, B.W., Wadsley, J., & Couchman, H.M.P., 2016, MNRAS, 463, 1431

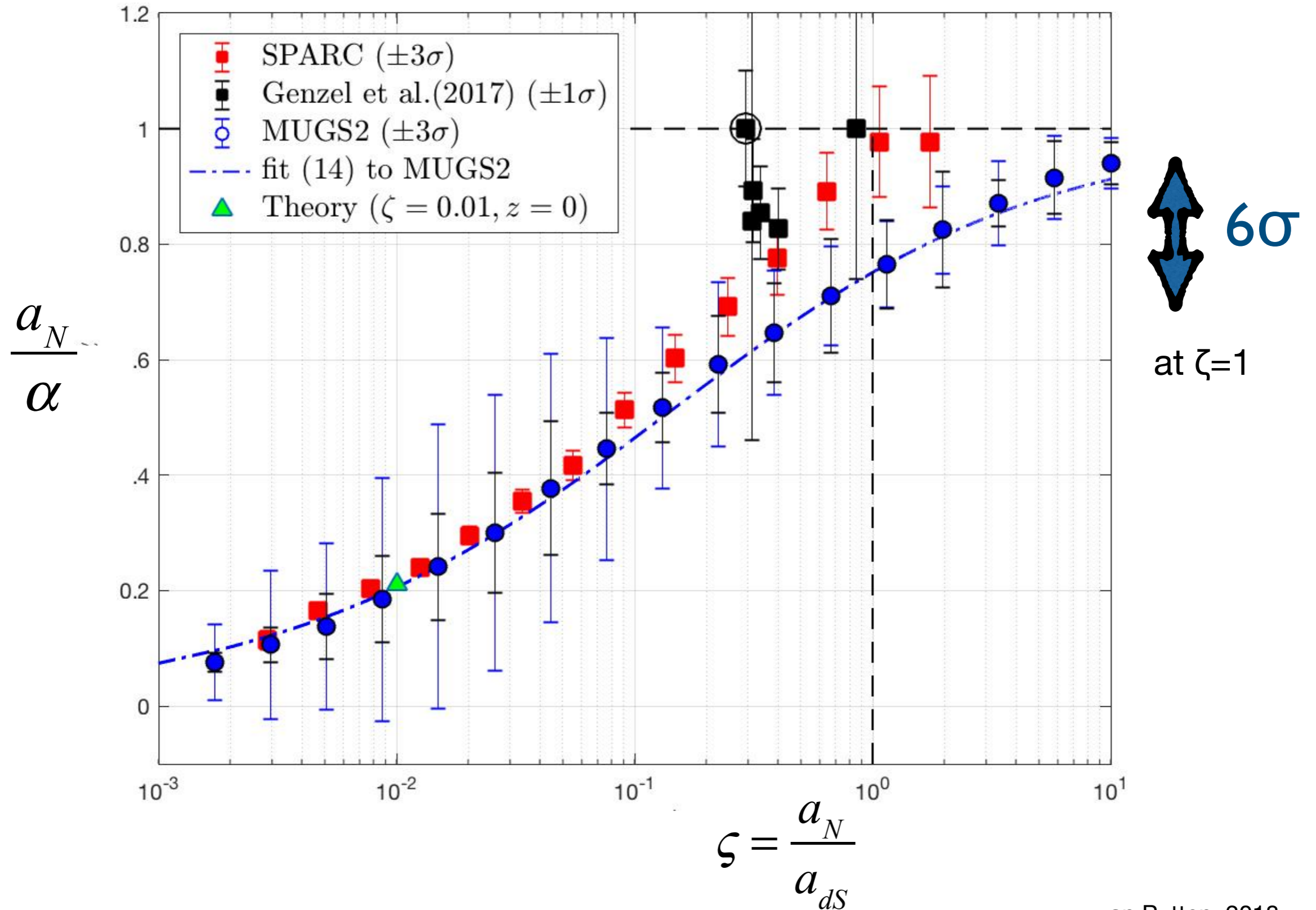
Keller, B.W., & Wadsley, J.W., 2017, ApJ, 835, L17

Self-similar galaxy dynamics

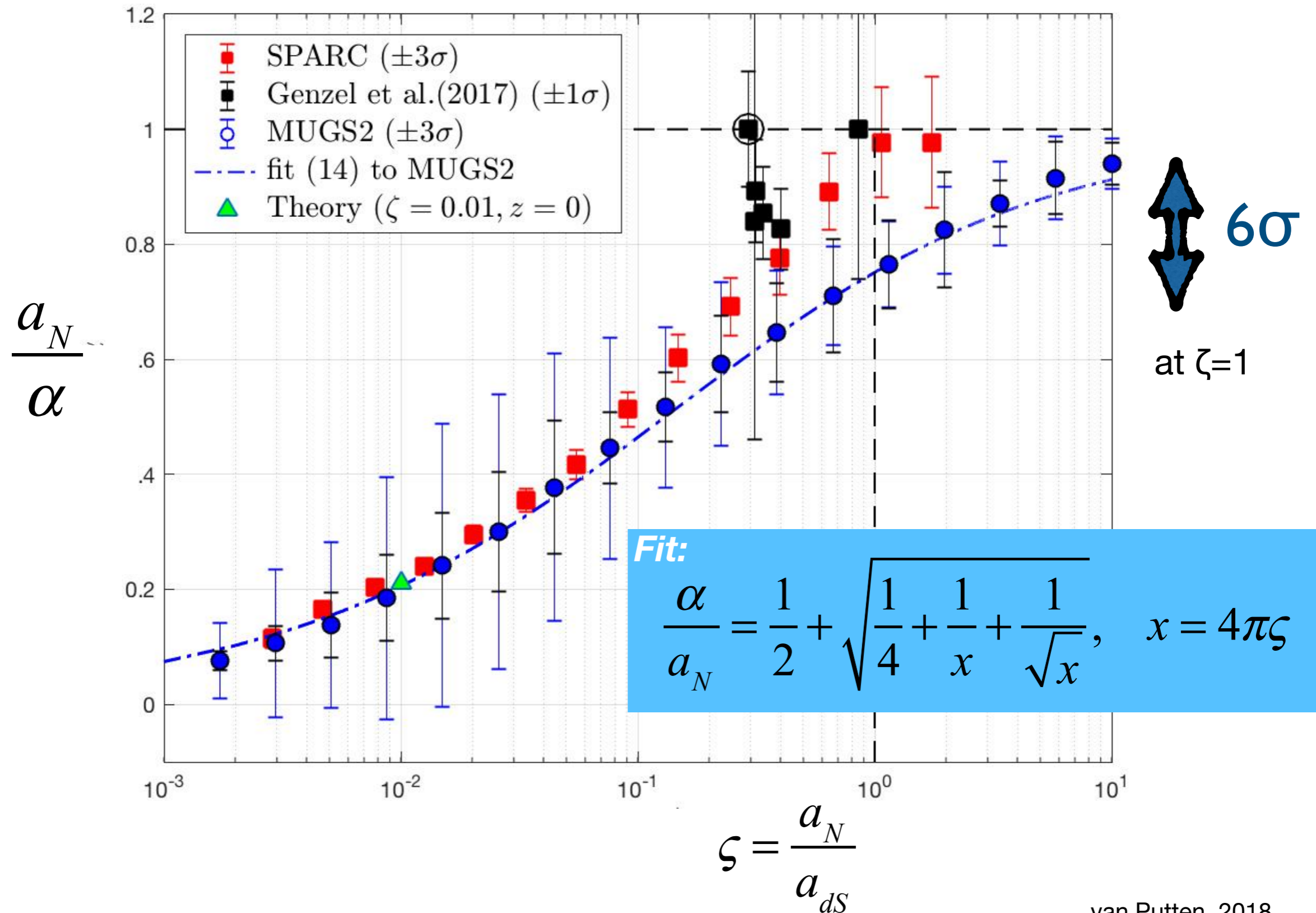
tracing background cosmology

van Putten, 2018, MNRAS 481 L26

Λ CDM/MUGS2 vs SPARC

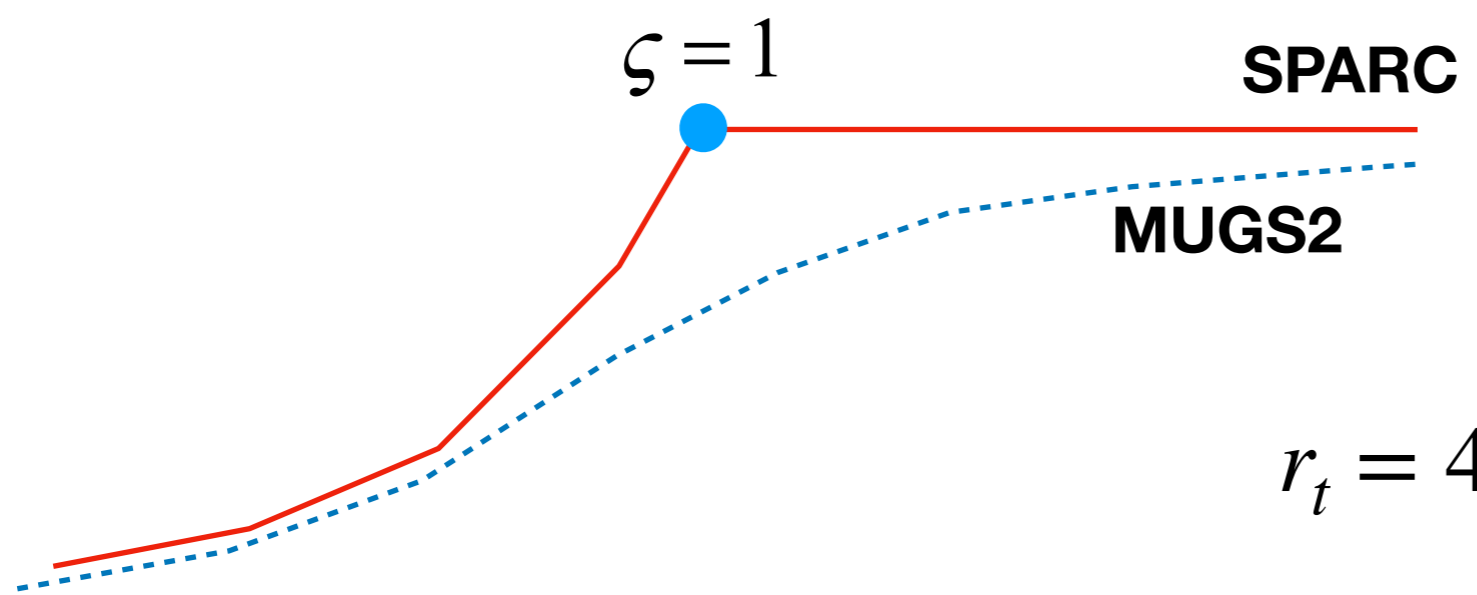


Λ CDM/MUGS2 vs SPARC

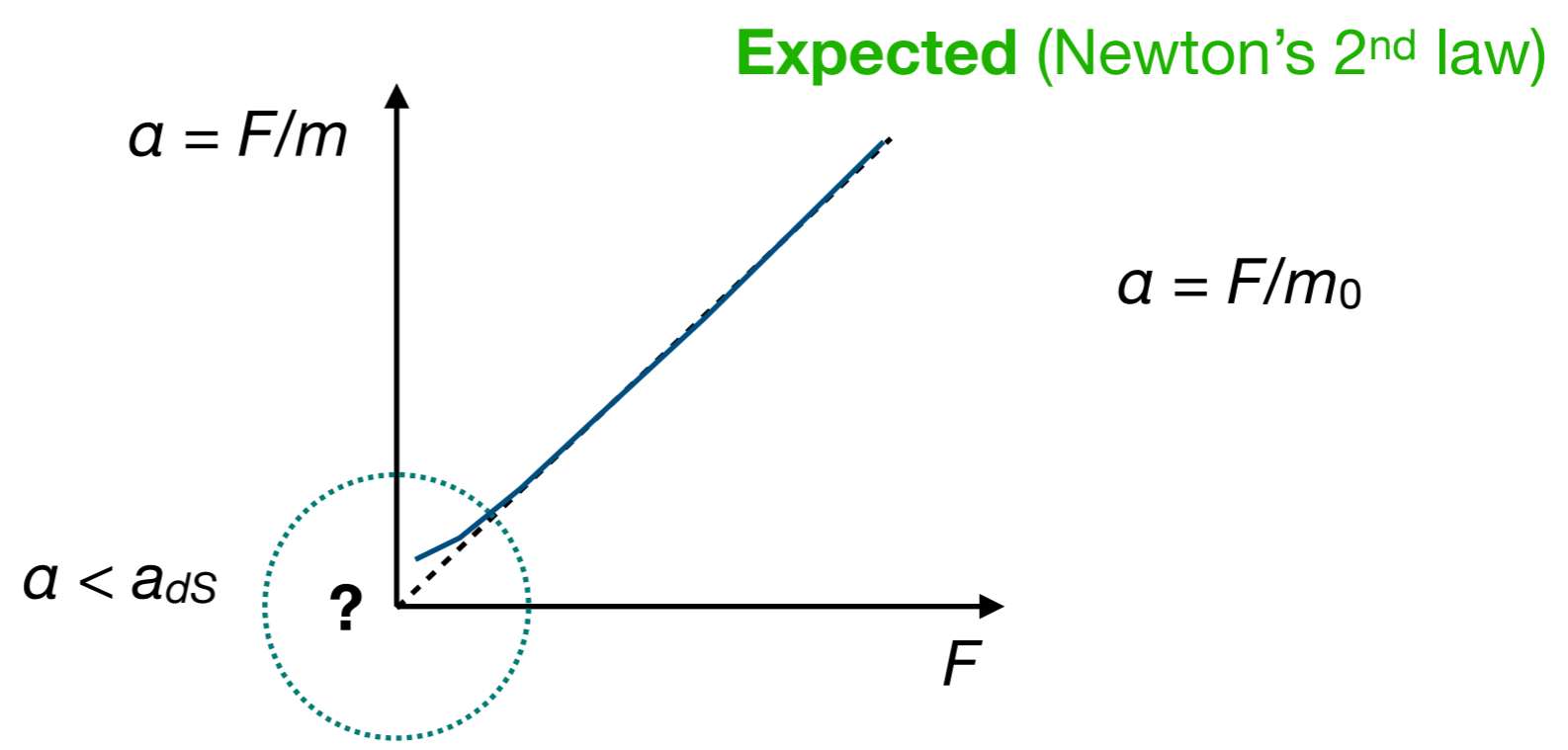


van Putten, 2018,
MNRAS, 481, L26

C⁰ Galaxy dynamics

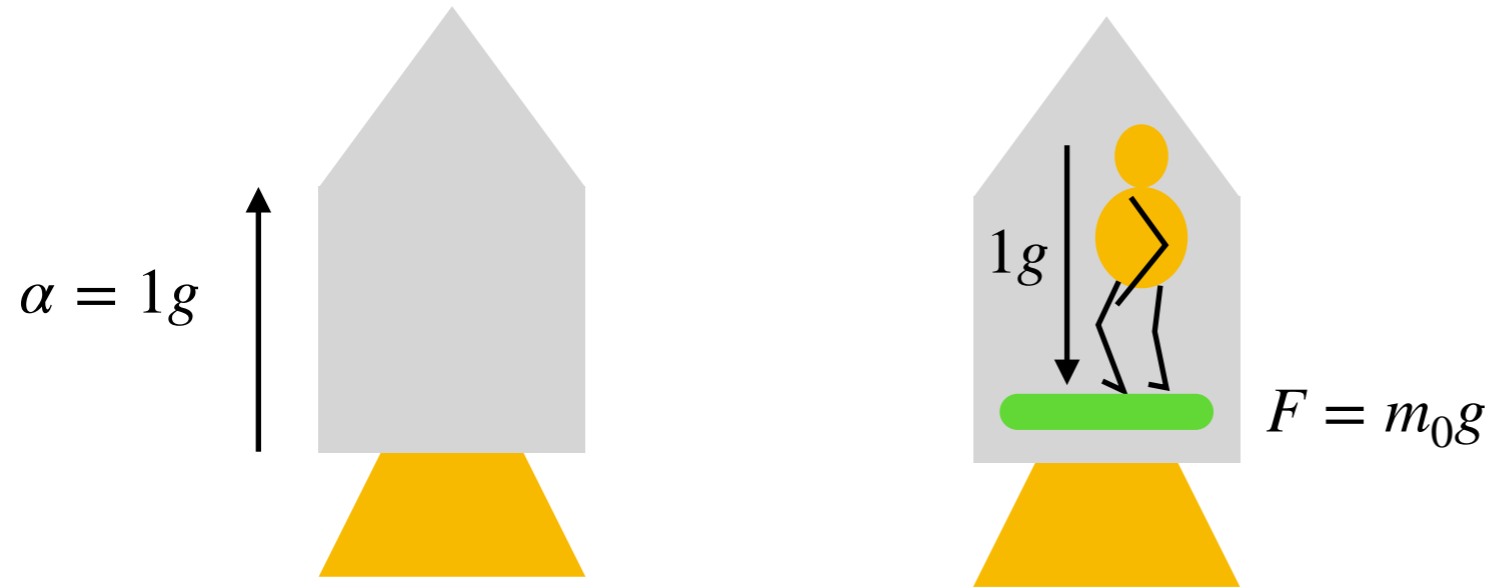
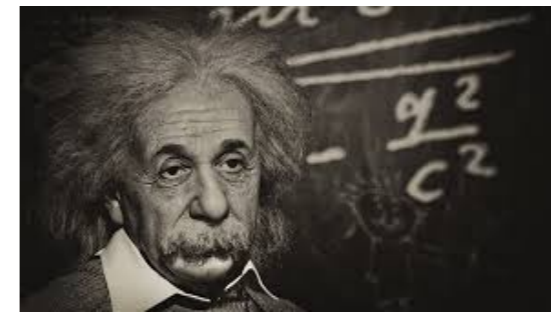


$$r_t = 4.7 M_{11}^{1/2} (H_0/H)^{1/2} \text{kpc}$$

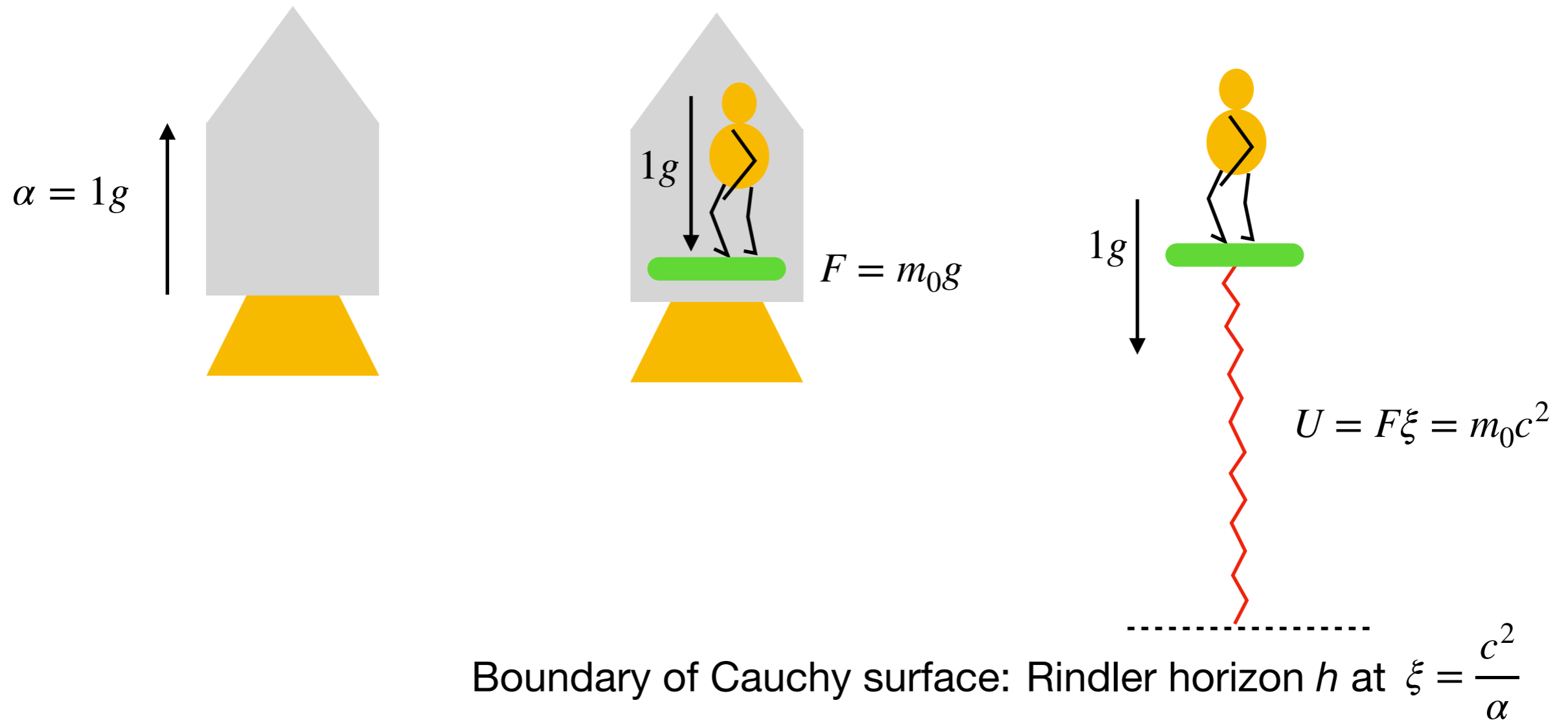
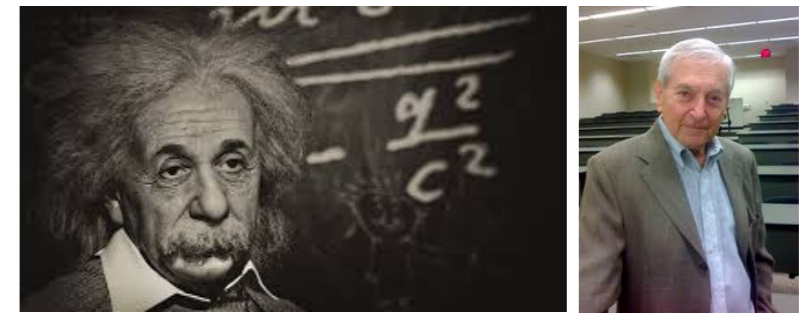


6 σ gap: Problem of inertia... ?

Equivalence Principle



Extended Equivalence Principle

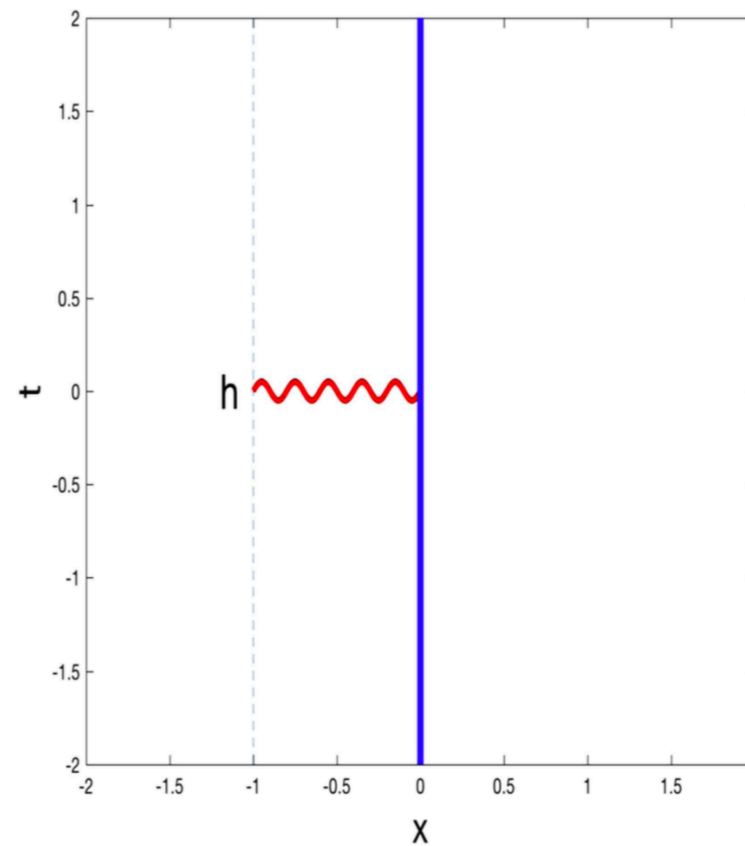
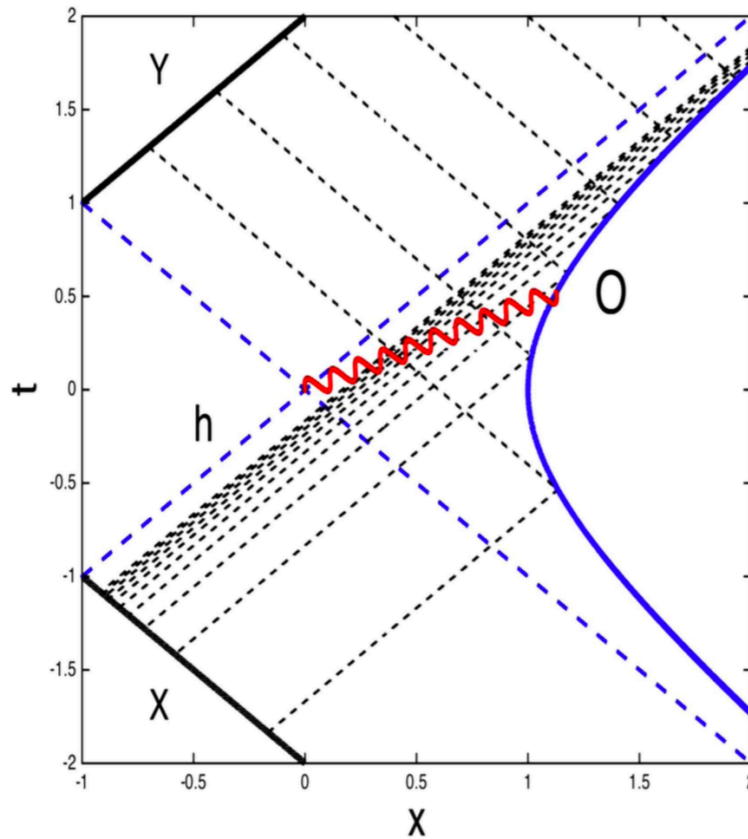


Newtonian inertia = Gravitational binding energy to h

Origin inertia is nonlocal

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Origin: entanglement entropy at Unruh temperature h set by α .

On a cosmological background with cosmological horizon \mathcal{H} at Hubble radius

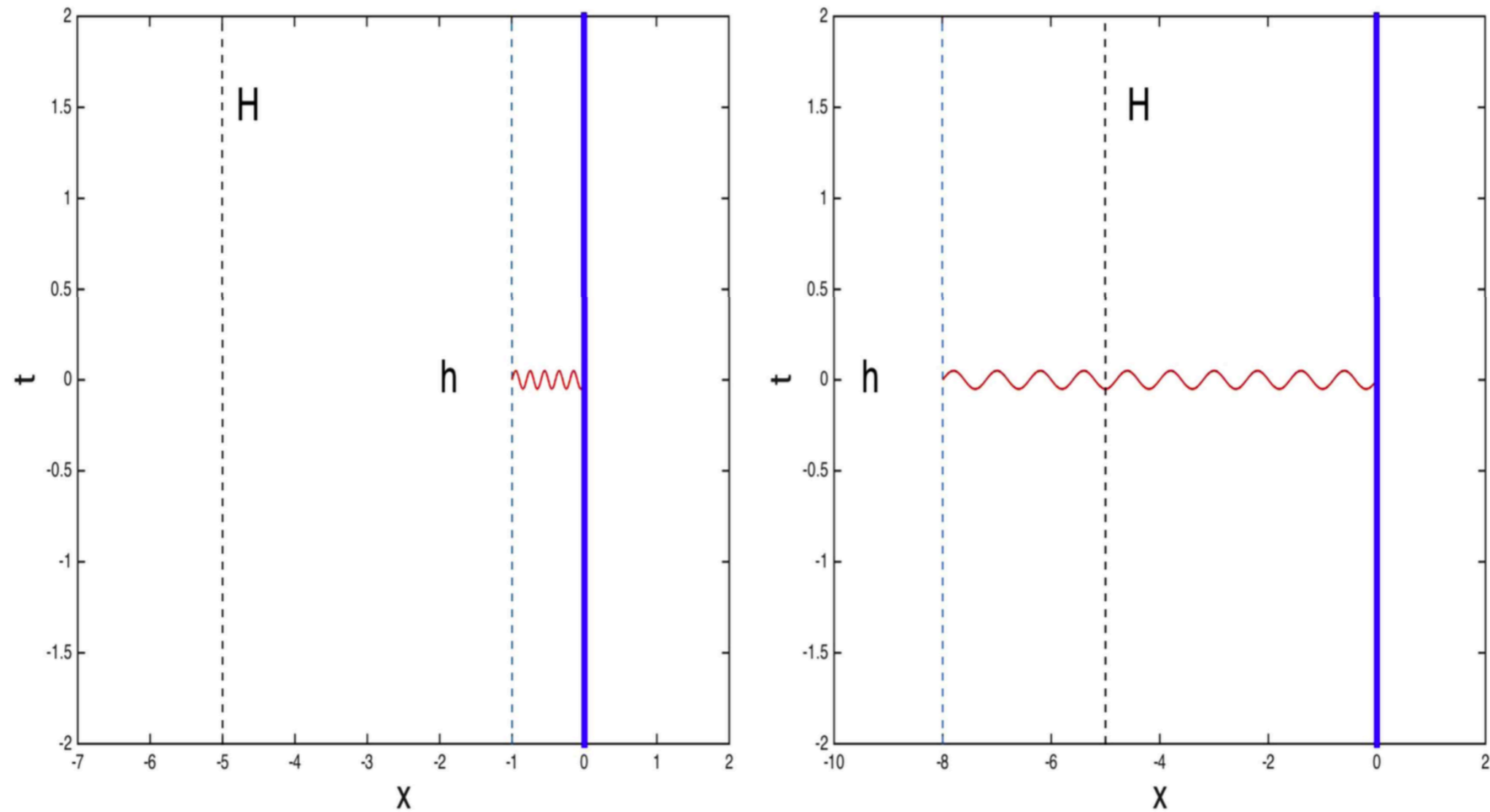
$$R_H = \frac{c^2}{a_{dS}} = \frac{c}{H}$$

h and \mathcal{H} are both apparent horizons ...

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\mathcal{H} -constraint on inertia

U to boundary of Cauchy surface: up to h or \mathcal{H} , whichever is more nearby (causality):

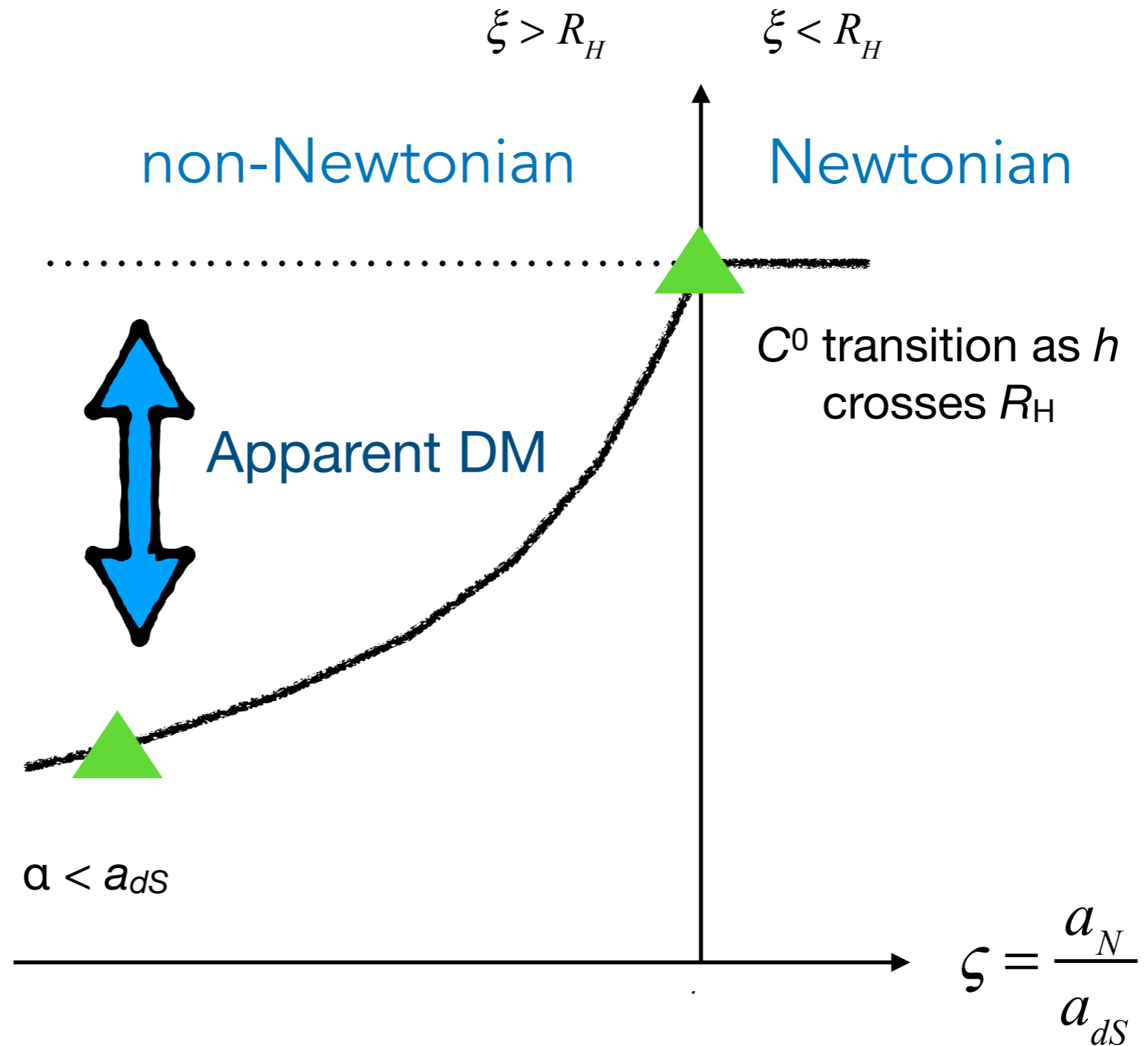


C^0 transition in U as α crosses a_{ds}

C⁰ galaxy dynamics

$$m\alpha = F_N = m_0 a_N :$$

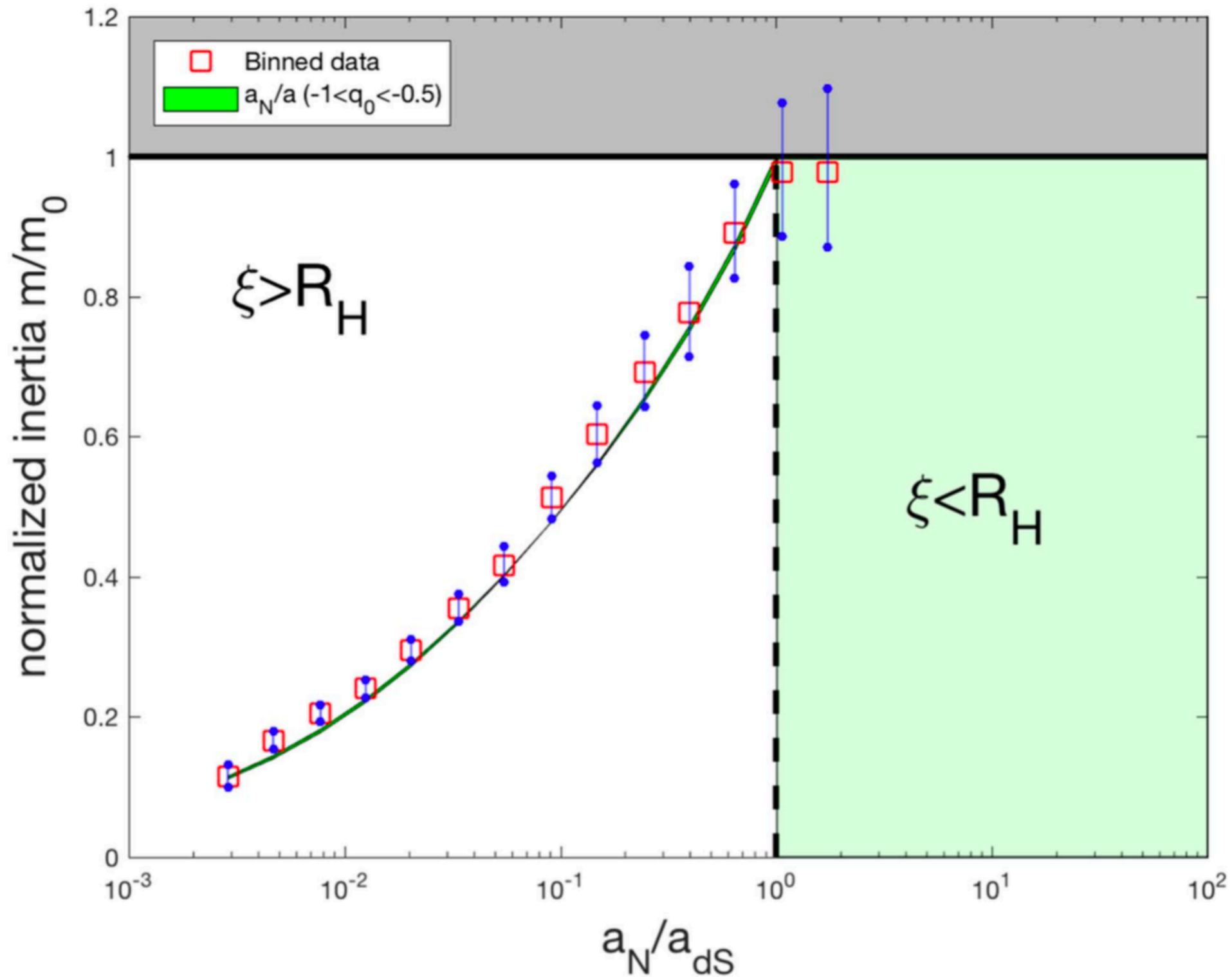
$$\frac{m}{m_0} = \frac{a_N}{\alpha}$$



Same E_k and U_N : invariant Lagrangian and Hamiltonian

Confrontation with SPARC

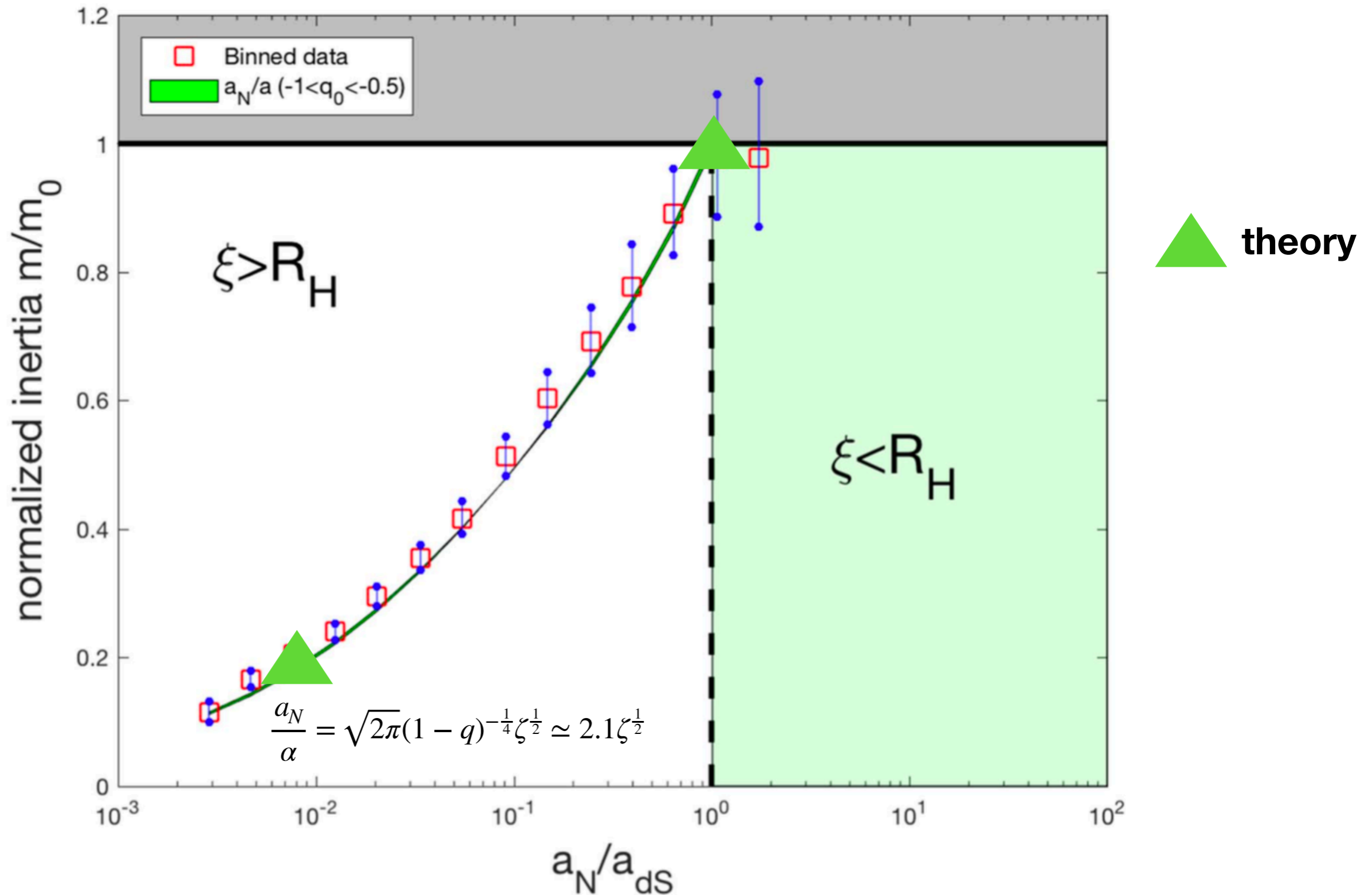
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van Putten, 2017, ApJ, 837, 22

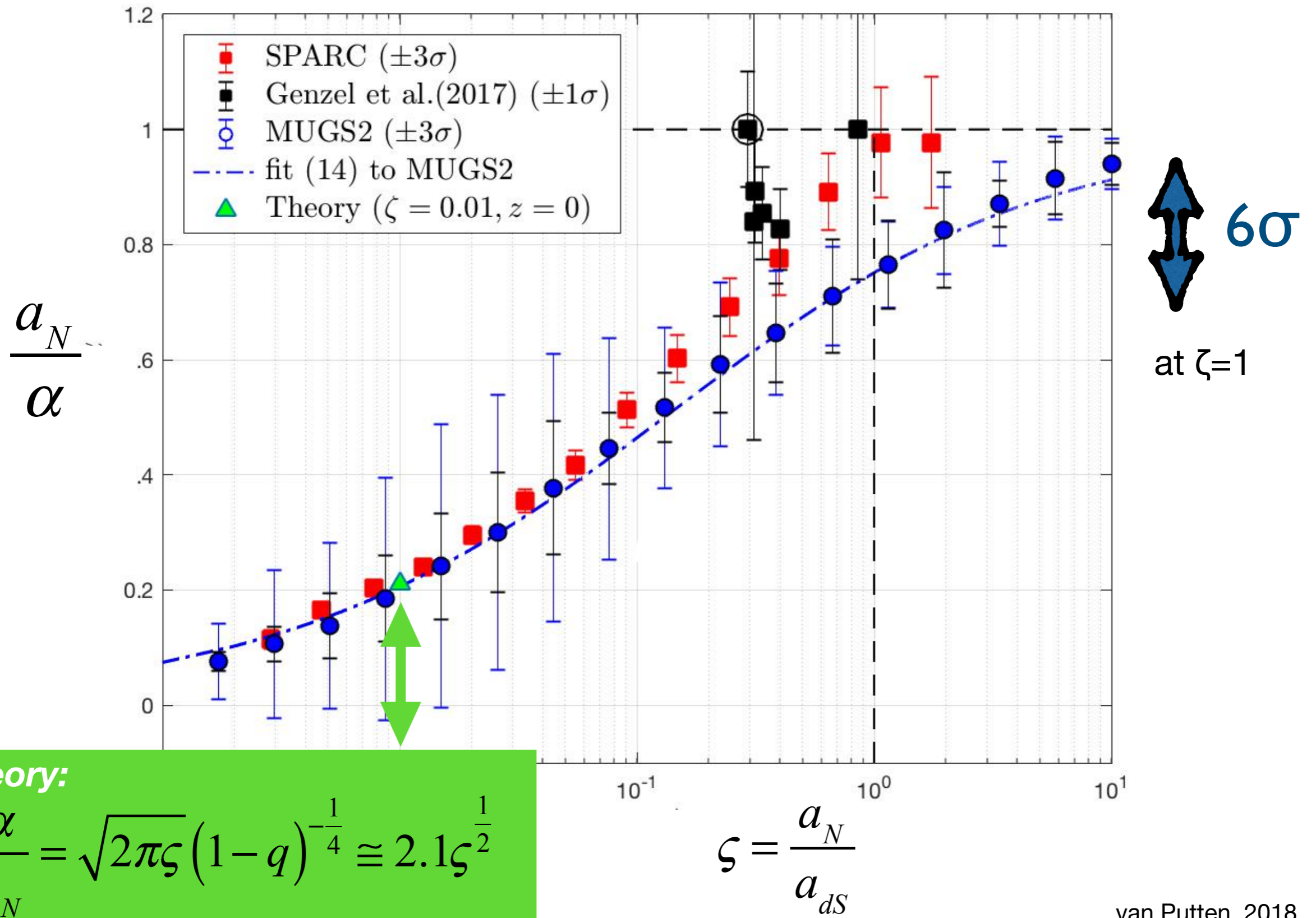
Confrontation with SPARC

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van Putten, 2017, ApJ, 837, 22

Λ CDM/MUGS2 vs SPARC



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Conclusions

Galactic evolution in a cosmological vacuum with a running de Sitter scale of acceleration $a_{\text{dS}} = cH$:

SPARC galaxies appear C^0 , effectively self-similar in $\zeta = a_{\text{N}}/a_{\text{dS}}$

6σ gap Λ CDM/MUGS2-SPARC about $\zeta=1$

Confrontation of Newton's 2nd law with SPARC

$m < m_0$ ($\zeta < 1$), $U = mc^2$ drops below Newtonian value $U = m_0c^2$ (h crosses \mathcal{H}).

No DM on galactic scales

$$\lambda \gg r_t = 4.7 M_{11}^{1/2} (H_0/H)^{1/2} \text{ kpc}$$

$m_{\text{DM}} \ll \text{mass of fuzzy DM}$

Expected tangible result from lab experiments:

accurate null-results

van Putten, 2015, MNRAS, 450, L48;
2018, MNRAS, 481, L26

van Putten, 2017, ApJ, 837, 22; ApJ, 848, 28